

Lecture 07: **Filters**

Wenda Cao

Big Bear Solar Observatory New Jersey Institute of Technology





Big Bear Solar Observatory

Outline





- Spectroscopy: spectrographs vs filters
- □ Interference Filters
- Lyot Birefringent Filters
- □ Fabry-Perot Interferometers
- Application in Astronomical
 Observation

Textbook: The Fabry-Perot Interferometer: History, Theory, Practice and Application, J M Vaughan

Big Bear Solar Observatory

New Jersey's Science & Technology University



1. Imaging Spectroscopy

□ Spectrographs vs Filters



Filter observes y, x

Big Bear Solar Observatory

efficient by a factor of $N_x/N_\lambda \sim 100$.



Spectrograph Spectroscopy





Х

Need x, y, λ observations, can observe 2 simultaneously

Spectrograph observes x, λ Filter observes y, x



If the spatial/spectral resolution and throughputs are the same and $N_x = N_y = N_\lambda$, then the two methods have same photons/time. Generally, $N_\lambda \ll N_x$ and filter is more efficient by a factor of $N_x/N_\lambda \sim 100$.



Big Bear Solar Observatory



Filter Spectroscopy



Big Bear Solar Observatory

Spectrograph observes x, λ Filter observes y, x If the spatial/spectral resolution and throughputs are the same and $N_x = N_y = N_\lambda$, then the two methods have same photons/time. Generally, $N_\lambda \ll N_x$ and filter is more efficient by a factor of $N_x/N_\lambda \sim 100$.





2. Interference Filters

- Interference filters are multilayer thin-film devices. An optical filter consisting of multiple layers of evaporated coatings on a substrate, whose spectral properties are the result of wavelength interference rather than absorption.
 - □ Interference filter classes
 - □ Interference filter structure
 - □ Interference filter principal
 - □ Interference filter terminology
 - □ Choose a right interference filter







Big Bear Solar Observatory



Filter Classes

- Short Wavelength Pass: transmits visible light of lower wavelengths and block light with higher wavelengths.
- Long Wavelength Pass: allows light of longer wavelengths to pass through it and effectively block shorter wavelengths.
- Band Pass: transmit one particular region (or band) of light spectrum. It passes only a very narrow region of wavelengths and blocks a majority of light incident upon the filter surface.
- □ **Sharp Cutting:** eliminates spectral regions, such as the infrared, "hot rejector".
- Broad Band: transmit one particular region (or band) of light spectrum. It usually has rather broad transmission characteristics and passes a significant number of wavelengths.





Big Bear Solar Observatory



Interference Filter Structure

- Interference filters are designed to provide constructive or destructive interference of light by taking advantage of the refraction of light through different materials.
- Glass substrates
- Multilayer thin-film coatings are applied to substrates.
- □ Single cavity bandpass filter
 - Spacer: the gap between the reflecting surfaces is a thin film of dielectric material, with a thickness of one-half wave at the desired peak transmission wavelength.
 - Reflection layers: consist of several film layers, each of which is a quarterwave thick.
- Multi-layer blocking filter
- Optical epoxy and protective metal ring







Big Bear Solar Observatory

New Jersey's Science & Technology University

How does it work ?

- □ Spacer: the gap between the reflecting surfaces is a thin film of dielectric material, with a thickness of one-half wave at the desired peak transmission wavelength. ($d = \lambda/2$)
- □ Start from a Fabry-Perot etalon ...

 $2AB - CD = 2d \cos \alpha$ $m\lambda = 2d \cos \alpha$

□ Constructive interference occurs when

 $m = 1, 2, 3 \dots$

□ Zero transmission occurs when m = 1/2, 3/2, 5/2...

 \Box Consider $d = \lambda/2$ and the normal incidence

$$\lambda_{trans} = \frac{2d\cos\alpha}{m} = 2\frac{\lambda_0}{2}\cos0^\circ = \lambda_0$$

 $\Box \quad How about the light of \quad \lambda \neq \lambda_0 \quad ?$





Specer at helf wavelength for the destroit wavelength, or a multiple of that.



New Jersey's Science & Technology University



How does it work ?

□ Reflection layers: consist of several film layers, each of which is a quarterwave thick $(d = \lambda/4)$.

 $m\lambda = 2d\cos\alpha$

□ Total reflectance on interface R for the normal incident light $R = \left(\frac{n_2 - n_1}{n_2 + n_1}\right)^2$

Consider $n_1 = 1$ and $n_2 = 1.5$ and 16 surfaces, R = ?

- **Constructive interference occurs when** m = 1, 2, 3...
- \Box Zero reflection occurs when m = 1/2, 3/2, 5/2...
- Then $m\lambda = 2d\cos\alpha = 2\frac{\lambda}{4}\cos0^\circ = \frac{\lambda}{2}$ m = 1/2
- So, zero reflection occurs and the light pass through the reflective surfaces if the reflected beam are in phase.



Refraction and Reflection in Optical Coatings







Principal Summary

- □ Reflection layers: consist of several film layers, each of which is a quarter-wave thick ($d = \lambda/4$).
- With the reflected rays being effectively cancelled, a thin film of quarter-wave thickness functions as an anti-reflection optical coating.
- □ Spacer: the gap between the reflecting surfaces is a thin film of dielectric material, with a thickness of one-half wave at the desired peak transmission wavelength. ($d = \lambda/2$)
- □ The gap in spacer determines which wavelengths destructively interfere and which wavelengths are in phase and will ultimately pass through the coatings.
- This principle strongly attenuates the transmitted intensity of light at wavelengths that are higher or lower than the wavelength of interest.



Reflection and Transmission by Interference Filters





Big Bear Solar Observatory





More Detail about Structure

- Spacer is the gap between the reflecting surfaces, which is a thin film of dielectric material.
- On either side of this gap are the two reflecting layers, which actually consist of several film layers.
- This sandwich of quarter-wave layers is made up of an alternating pattern of high and low index material, usually ZnS (n=2.35) and cryolite (n=1.35). Together, they are called a stack.
- □ The number of layers in the stack is adjusted to tailor the width of the bandpass.
- To sharpen cutoff, it is common practice that several cavities are layered sequentially into a multicavity filter, which dramatically reduces the transmission of out-of-band wavelengths.





Big Bear Solar Observatory



Interference Filter Anatomy





Big Bear Solar Observatory

Terminology



- Bandpass: the range (or band) of wavelengths passed by a wavelength-selective optic.
- Blocking: the degree of light attenuation at wavelengths outside the passband of filter.
- □ **Center Wavelength (CWL):** the wavelength at the midpoint of the half power bandwidth (FWHM).
- □ **Full-Width Half-Maximum (FWHM):** the width of the bandpass at one-half of the maximum transmission.
- **Peak Transmission:** the maximum percentage transmission within the passband.
- □ *Filter Cavity:* An optical "sandwich" of two partially reflective substrate layers separated by an evaporated coating which forms the dielectric spacer layer.



Interference Filter Characteristics and Nomenclature

Big Bear Solar Observatory

New Jersey's Science & Technology University

Performance

- Transmitted distortion: the distortion of a plane wavefront passing through the filter, and is also measured in fractions or multiples of a wavelength.
- Wedge: angular deviation from parallelism between the outer filter surfaces, which is measured in aresecond or are-minutes of the deviation angle.
- Angle shift: the wavelength of CWL at small angle *φ* from normal incidence is

$$\lambda = \lambda_0 \sqrt{1 - \left(\frac{n_0}{n_e}\right)^2 \sin^2 \phi}$$

where $n_0 = 1$ in air, n_e is refractive index of spacer material.

- Temperature: an interference filter is slightly temperature dependent, causing transmission spectrum shifts slightly to longer wavelengths with increasing temperature.
- Orientation: the shiniest side toward the source.

Big Bear Solar Observatory

Wavefront Distortion by Optical Surfaces Reflected Distortion Transmitted Distortion Wedge Deviation (a) Reflected Incident Optical Transmitted (c) Wavefront Element Wavefront Transmitted Wavefront Transmitted Wavefront Transmitted Transmitted Wavefront Transmitted Transmitted Transmitted Wavefront Transmitted Tr

Temperature Dependence of Peak Transmittance

	Temperature
Wavelength	Coefficient of Shift
(nm)	(nm per °C)
400	0.016
476	0.019
508	0.020
530	0.021
557	0.021
608	0.023
630	0.023
643	0.024
710	0.026
820	0.027

New Jersey's Science & Technology University



Select a Right Filter



Quotation

For:	Wenda Cao		From:	Kyle Bushong		
	Big Bear Solar Observatory		Phone:	978-692-7513 x3658		
Address:	New Jersey Institute of Technology		Fax:	978-692-7443		
	40386 North Shore Lane		E-Mail:	kyle_bushong@beminc.com		
	Big Bear City, CA 92314-9672					
Phone:						
Fax:						
E-Mail:	wcao@email.noao.edu	See Attached				

Description Unit Price Total Price Item Qtv CWL: 430.5 +0.2/-0nm 2 \$ 3,720.00 1 \$ 7,440.00 FWHM: 1.0 ± 0.2nm Peak %T: > 60% Out of band blocking: E-5 from 200-1200nm Transmitted wavefront: ≤ 1/4wv PV per 1" CA @ 632.8nm Size: 50mm diameter nominal Clear aperture: 45mm minimum Thickness: unspecified Angle of incidence: 0° Operating temp: -10 to 30°C Surface quality: 60/40 per MIL-C-48497A Multiple substrate air-spaced construction, laminated into AI ring Air gap ~0.5-0.6mm Parts will be tested in air @ ~23°C with a collimated beam 2 Estimated shipping/insurance charge 1 \$ 65.00 65.00 s Shipping: ~\$25 Insurance: ~\$40 Minimum order value: \$5000.00 USD Deliverable Data: In-band evaluation per filter, confirm peak T%, CWL, FWHM Out-of-band blocking from 200-1200nm on 1 filter per batch ONLY One scan per filter to confirm TWF, ≤ 1/4wv per 1" clear ap. Air-space worst case: add TWF of both components Pricing above does NOT include freight. All freight will be prepaid & N KNOWLEDGE added to invoice.



Select a Right Filter



Prepar	ed for:	Wenda C NJIT USA	ao			Qu Issu Qua Pay F.C	otation ue Date: ote Valid yment To).B.: All pr	n # : dity: erms: <i>ices a</i> l	283 5/12/2 90 day Pendir Salem <i>re quo</i>	07 010 /s from (ng , NH ted in (date U.S. d	lollars	
	Fax: Phone: e-mail:	wcao@er	nail.noao.edu			Pre Re.	pared b	y: Ph phi RF	illip Clai il.clark@ Q dateo	rk 9andove 1 May 1(ercorp.c Oth	com	
<u>Item #</u> 1	Descri Wave Bandy Trans Blocki	ption length: width: mission: ing: rustion:	430.5 nm +0.1/-0 0.5 nm ±0.1 40% 1 x 10-4 avg. X-Ray to 1200r	۱m		<u>(</u>	2 <u>tv</u> 1 2	Unit pri \$3,995. \$3,295.	<u>ce</u> 00 00	Extende Item prie \$3,995.0 \$6,590.0	e d <u>ce</u> 00	<u>Dellverv</u> 3-4 weeks ARO 3-4 weeks ARO	2
	Size: Thicki Clear Temp	ness: Aperture: . (° C.):	50.0 mm +0/-0.25 7.0 mm maximum 45.0 mm	SC DIC	RATCH: / 5: A small Scratch or Dig Number #	CH: Any marking or tearing of the part surface. small rough spot on the part surface similar to a pit in appearance. A bubble atch or Maximum Maximum Dig Dig or Bubble wimber Scratch Width Or Bubble Diameter Separation Distance # mm inch mm inch mm inch							considered
	Opera Subst TWF: Parall Scrate Polari	ating Angle: rate Mat'l: elism: ch/Dig: zation:	0 BK7 & Filter Glass 1/4 wave per inch or better 30 arc seconds or better 60/40 Bandom		120 80 60 50 40 30	0.12 0.08 0.06 0.05 0.04 0.03	0.0047 0.0031 0.0024 0.0020 0.0016 0.0012	1.20 0.80 0.60 0.50 0.40 0.30	0.0473 0.0315 0.0236 0.0196 0.0158 0.0118	20 20 20 20 20 20 20 20	0.787 0.787 0.787 0.787 0.787 0.787 0.787		



Effective Index: Comments:

1. Exterior surfaces will be A/R coated 2. Filter will be mounted in a ring

1.45

20 0.02 0.0008 0.20 0.0079 20 0.787 15 0.015 0.0006 0.15 0.0059 20 0.787 10 0.010 0.0004 0.10 0.0039 1.0 0.040 5 0.005 0.0002 0.05 0.0020 1.0 0.040 3 0.003 0.00012 0.03 0.0012 1.0 0.040

EDGE

3. Lyot Filters





Bernard Lyot 1897-1952

- A Lyot filter, named for its inventor Bernard Lyot, is a type of optical filter that uses <u>birefringence</u> to produce a <u>narrow passband</u> of transmitted wavelengths.
- Lyot filters are often used in astronomy, particularly for solar astronomy.





Big Bear Solar Observatory

Birefringence

- Birefringence, or double refraction, is the decomposition of a ray of light into two rays when it passes through certain anisotropic material (birefringent crystal), such as crystals of calcite.
- When a beam of light is incident on a birefringent crystal, the waves are split upon entry into orthogonal polarized components: ordinary and extraordinary.
- o and e components travel through the molecular lattice along different pathways, depending on their orientation with respect to the crystalline optical axis.
- Light passing through a birefringent crystal
- □ **Parallel entry:** o and e wavefront coincide in amplitude, phase, and trajectory during their journey in the crystal.
- □ **Oblique entry:** o and e diverge and follow different pathways, and **o wave travels faster then e wave**.
- Perpendicular entry: divergence between o and e is eliminated, but o wave still travels at a higher speed than does e wave.



Separation of Light Waves by a Birefringent Crystal





 $n_o \neq n_e$



Big Bear Solar Observatory

Birefringence and Interference

Perpendicular entry: the propagation speed of o and e wave differ. Birefringent index of a crystal is defined as

$$\mu = n_e - n_o$$

o and e wave travel through a crystal of thickness d with a phase delay

$$\delta = \frac{2\pi(\Delta OPL)}{\lambda} = \frac{2\pi d(n_e - n_o)}{\lambda} = \frac{2\pi \mu d}{\lambda}$$

Consider a birefringent crystal of a thickness of d, which is placed between two linear polarizers with the same polarization direction. Assume the optical axis of the crystal is 45 with respect to the polarization directions, then the transmitted light is given by

$$T = \cos^2(\frac{\delta}{2}) = \cos^2(\frac{\mu d}{\lambda}\pi) = \cos^2(\sigma\pi)$$

Separation of Light Waves by a Birefringent Crystal





Big Bear Solar Observatory

New Jersey's Science & Technology University



Transmission Profiles



$$T_1 = \cos^2(\frac{\delta_1}{2}) = \cos^2(\frac{\mu d_1}{\lambda}\pi) = \cos^2(\sigma_1\pi)$$

□ What does the transmission profile look like if $d_2 = 2d_1$?

$$T_2 = \cos^2(\frac{\delta_2}{2}) = \cos^2(\frac{\mu d_2}{\lambda}\pi) = \cos^2(\sigma_2\pi) = \cos^2(2\sigma_1\pi)$$



000





FWHM and FSR



Full Width at Half Maximum (FWHM): is determined by the thickness of the thickest stage d_{thick}.

$$\Delta \lambda_{FWHM} = \frac{\lambda^2}{2 \mu d_{thick}}$$

Free Spectral Range (FSR): is determined by the thickness of the thinnest stage d_{thin}.







Big Bear Solar Observatory

New Jersey's Science & Technology University

Lyot Filter Tuning



- □ Wavelength tuning is a critical feature to calibration, fabrication and operation.
- **Each** stage needs its individual tuning system.
- □ A quarter waveplate, which follows the crystal to be 45° with respect to the optical axis, is followed by a rotating polarizer or a rotating half waveplate.





BBSO NIR Lyot Filter

- Working Wavelength:
- **Clear Aperture:**
- Passband FWHM: 2.5 Å
- Tunable Range:
- Peak Transmission: \succ
- Internal Structure:
- Minimum tunable step: 0.01 Å

Fe I 1.5648 & 1.5652 μm ~ 37 mm ± 7 Å ~ 8 % for non-polarized light 4-module Thermal Controller: $35.000 \pm 0.005^{\circ}$ C







Design Requirement

Big Bear Solar Observatory THE EDGE IN KNOWLEDGE



Optical and Mechanical Design







Big Bear Solar Observatory

System Calibration





Big Bear Solar Observatory

New Jersey's Science & Technology University

System Calibration











Big Bear Solar Observatory



Transmission Profiles



New Jersey's Science & Technology University



Order Sorting Filter



Big Bear Solar Observatory



4. Fabry-Perot Interferometer





Fabry-Perot interferometer (FPI), also called Fabry-Perot etalon is made of two semi-reflecting plates of glass, parallel, producing an interference pattern.





Collection Ecole polytechnique





Big Bear Solar Observatory

How does a FPI work?

Start from a Fabry-Perot etalon ...

 $2AB - CD = 2\mu d\cos\alpha$

 $m\lambda = 2\mu d\cos\alpha$

Constructive interference occurs when

$$m = 1, 2, 3 \dots$$

directly
transmitted
+ = $\bigwedge_{\text{constructive}}_{\text{interference!}}$

Destructive interference occurs when

$$m = 1/2, 3/2, 5/2..$$







some light reflects four times...



Big Bear Solar Observatory

some light goes straight through...



New Jersey's Science & Technology University



Interference Fringe





When a FPI is illuminated by a monochromatic extended source.



□ When a FPI is illuminated by a polychromatic extended source ...

- $m\lambda = 2\mu d\cos\alpha$
- □ Constructive interference occurs when
 - $m = 1, 2, 3 \dots$

□ When incident angles are fixed, transmitted wavelength depends on the spacing of a FPI.



Big Bear Solar Observatory

Transmission Profile

□ The amplitude *E*_t(*m*) of the resultant electric vector of transmitted light is given by

$$E_{t}(m) = t_{1}^{+}t_{2}^{+}[1 + r_{1}^{-}r_{2}^{+}e^{i\varphi} + \dots + (r_{1}^{-}r_{2}^{+})^{m-1}e^{i(m-1)\varphi}]$$

$$E_{t} \to E_{t}(\infty) = t_{1}^{+}t_{2}^{+}/(1 - r_{1}^{-}r_{2}^{+}e^{i\varphi})$$

□ The energy transmission coefficient for the pair of surfaces:

$$I_{t} = E_{t}E_{t}^{*} = \left|t_{1}^{+}t_{2}^{+}\right|^{2} / (1 + \left|r_{1}^{-}r_{2}^{+}\right|^{2} - 2\left|r_{1}^{-}r_{2}^{+}\right|\cos\varphi)$$

$$= \frac{T^{2}}{(1-R)^{2}} \left(\frac{1}{1 + [4R/(1-R^{2})]\sin^{2}(\varphi/2)}\right)$$

$$= [T/(1-R)]^{2} [1 + F\sin^{2}(\varphi/2)]^{-1}$$



d, µ

$$\varphi = 2\pi (2\mu d\cos\theta)/\lambda$$

$$T = t^{+}t^{-}$$
$$R = (r^{+})^{2} = (r^{-})^{2}$$
$$R + T = 1$$

□ The energy reflection coefficient for the pair of surfaces:

$$I_R = F \sin^2 (\varphi/2) [1 + F \sin^2 (\varphi/2)]^{-1}$$





Transmission Profile

A + R + T = 1







New Jersey's Science & Technology University

Consider absorption

1.0



FWHM, FSR and Finesse

□ Full width at half maximum (FWHM):

 $fwhm = \frac{(1-R)\lambda^2}{2\pi\mu d\cos\theta\sqrt{R}}$

□ Free spectral range (FSR):

$$fsr = \frac{\lambda^2}{2\mu d\cos\theta} = \frac{\lambda}{m}$$

Finesse:

$$N_R = \frac{fsr}{fwhm} = \frac{\pi\sqrt{R}}{1-R}$$

□ *Resolving power:*

$$m\lambda = 2\mu d\cos\theta$$

$$\mathcal{R} = \frac{\lambda}{\Delta \lambda} = mN_R = \frac{2\mu d\cos\theta}{\lambda}N_R$$

Big Bear Solar Observatory

New Jersey's Science & Technology University



An Example











Big Bear Solar Observatory



FPI Characteristic Evaluation



Big Bear Solar Observatory

New Jersey's Science & Technology University



FPI Testing Results



Big Bear Solar Observatory

New Jersey's Science & Technology University

FPI Testing Results

Big Bear Solar Observatory



• FPI Control & Scan Step: ~ 0.00292 Å / Step



• Plate Flatness & Roughness:

Flatness ~ $\lambda / 306$ @ 1523 nm Roughness ~ $\lambda / 1842$ @ 1523 nm

• Peak Transmission:



15635 15640 15645 15650 15655 15660 15665 1567C Wavelength [Angstram]



~ 90.5 % @ 1523.1 nm

New Jersey's Science & Technology University



Application I – IRIM InfraRed Imaging Magnetograph





Big Bear Solar Observatory

Application II – Measuring Neutral Winds in the Upper Atmosphere





Big Bear Solar Observatory